

# Morphological Bulge-Disk Decompositions of Massive Galaxies at $z=1-3$ in CANDELS

Victoria Bruce

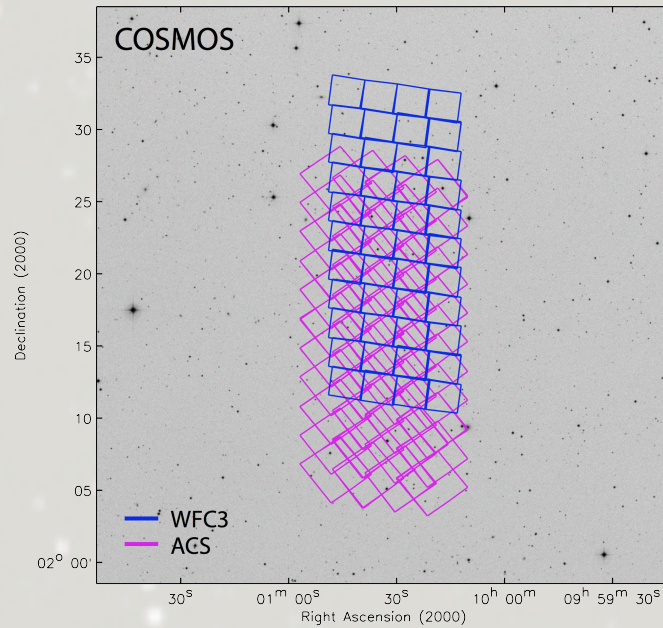
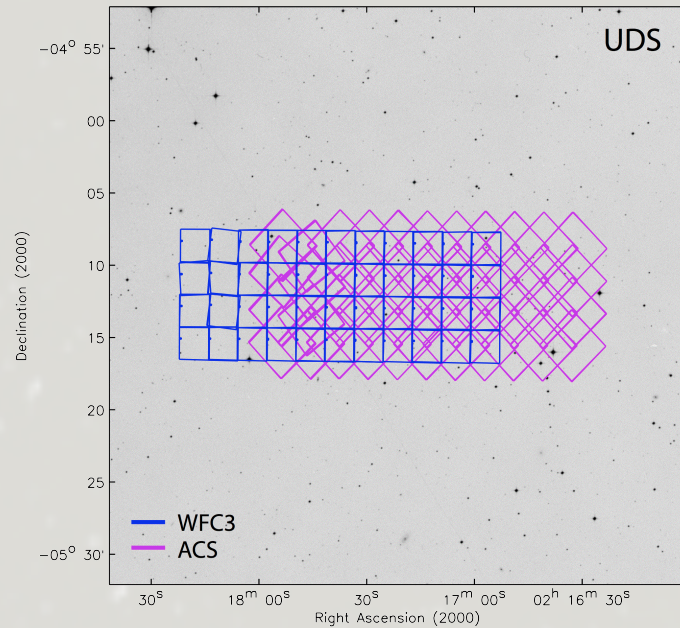
University of Edinburgh



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# Introduction

- How is star-formation quenched in massive galaxies and how is this linked to morphological transformations ?
- How can morphological bulge+disk decompositions help to answer these questions ?
- CANDELS:



Grogin et al. 2011

- Mass-selected, homogeneous sample of  $\sim 400$  objects within  $1 < z_{\text{phot}} < 3$  and with  $M_{\star} > 10^{11} M_{\odot}$

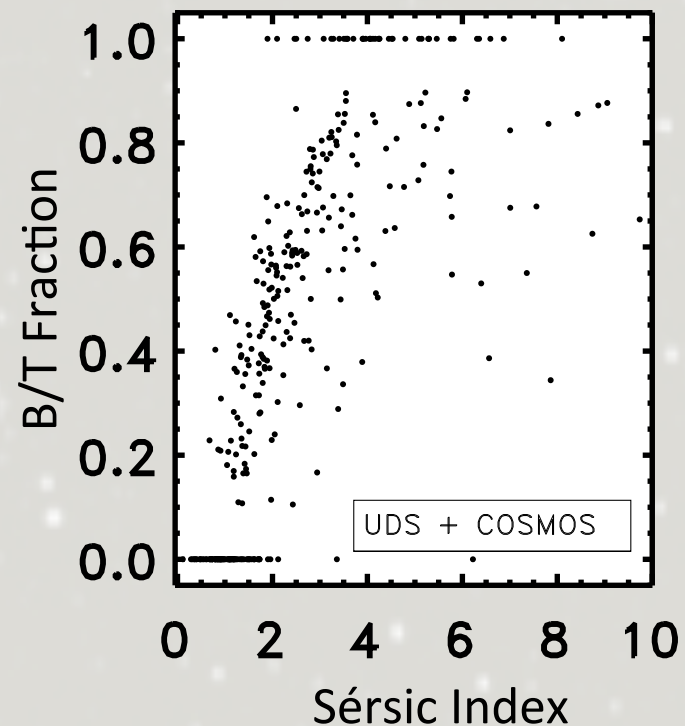
# Bulge+Disk Decompositions

Define 3 components : disk  $n=1$ , bulge  $n=4$ , PSF

Models:

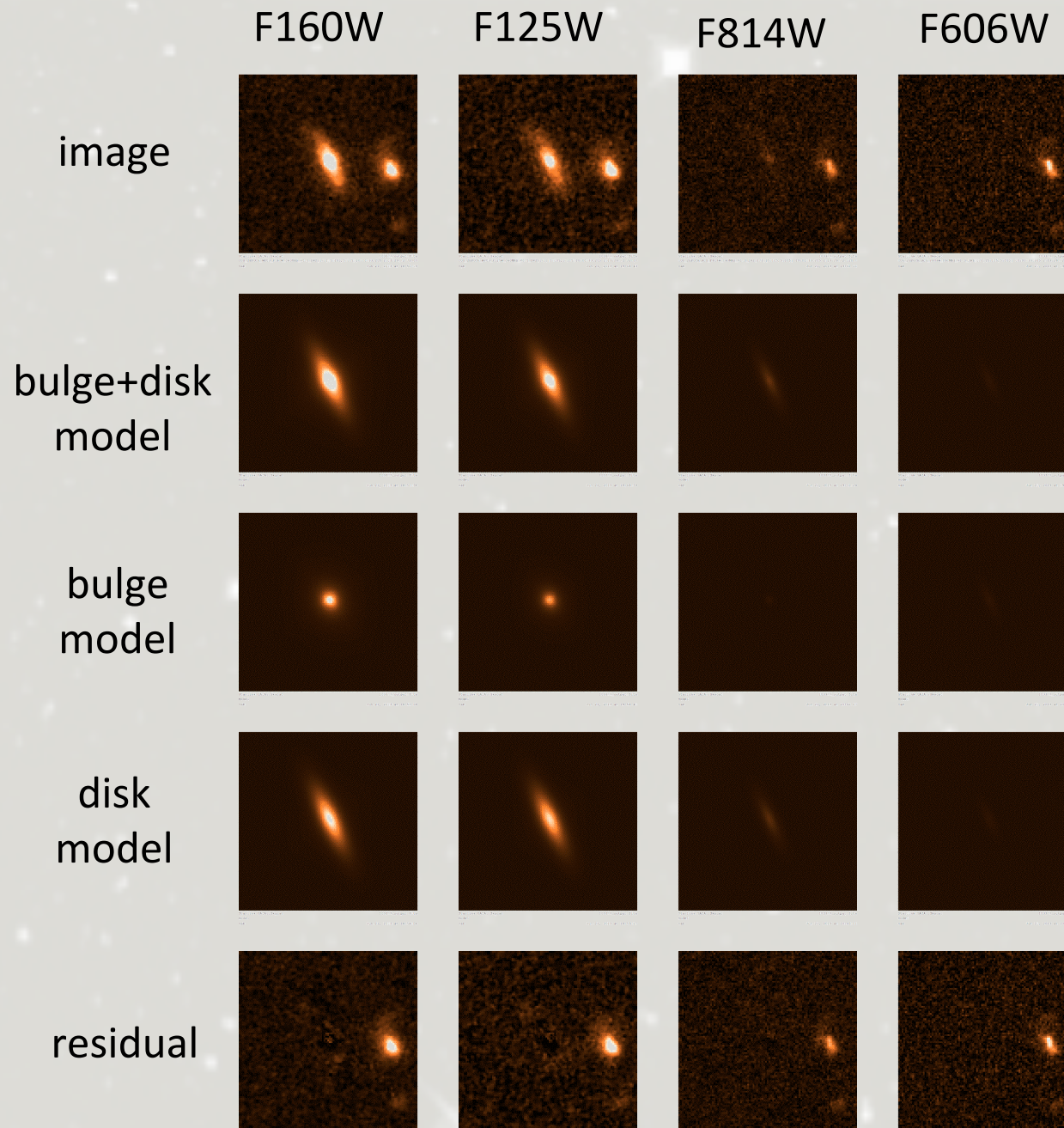
- bulge only
- disk only
- bulge + disk
- bulge + PSF
- disk + PSF
- bulge + disk + PSF

Relation between the bulge/total light fraction and the single component model Sérsic index from the F160W decompositions.





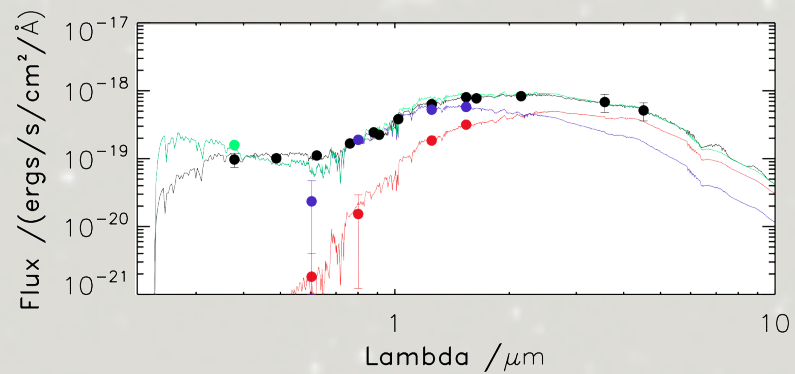
# Example Morphological Decomposition



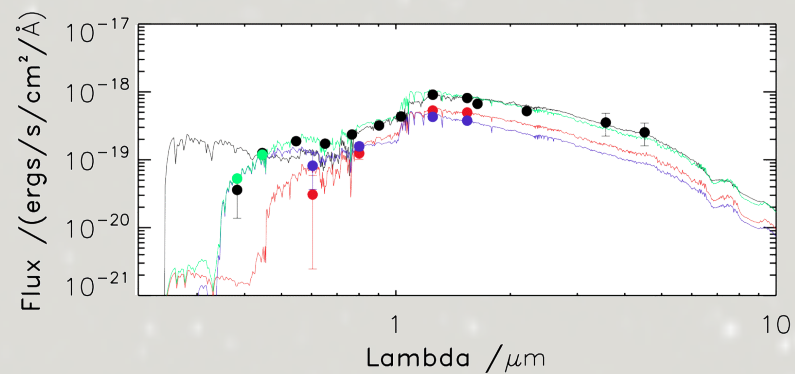


# Example SED Fitting

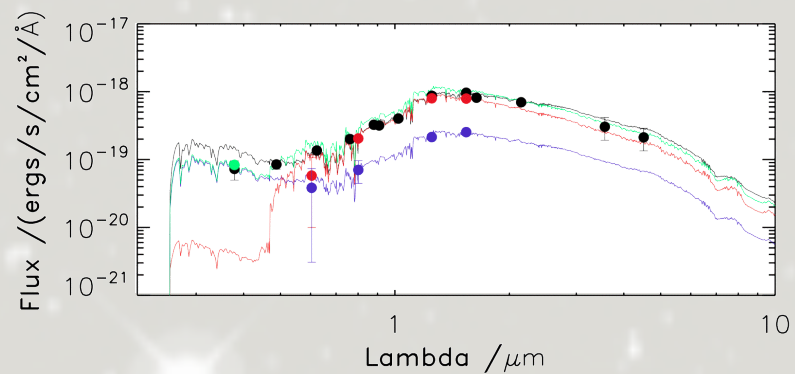
B/T < 0.5



B/T = 0.5



B/T > 0.5



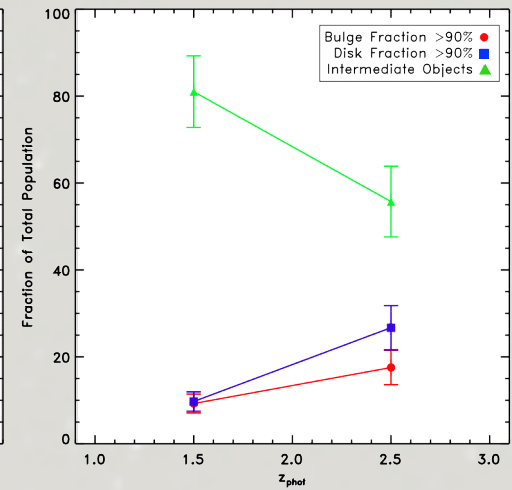
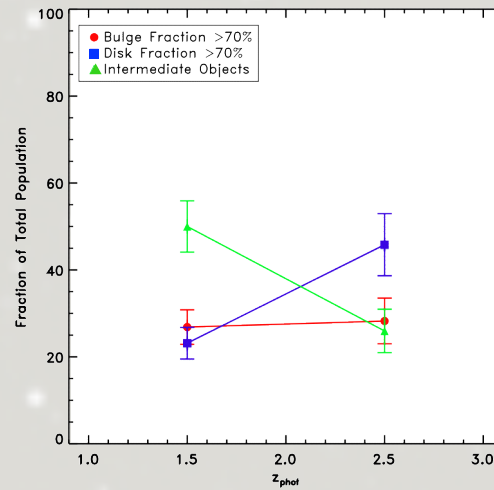
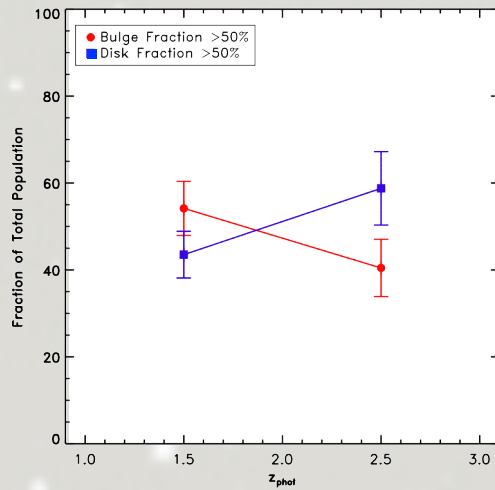
# Evolution of Morphological Fractions

B/T > 0.5  
D/T > 0.5

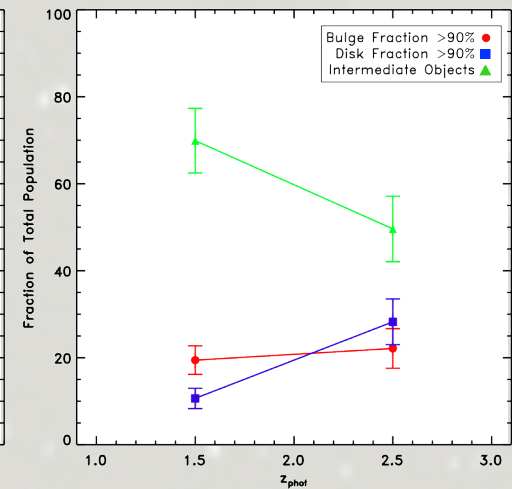
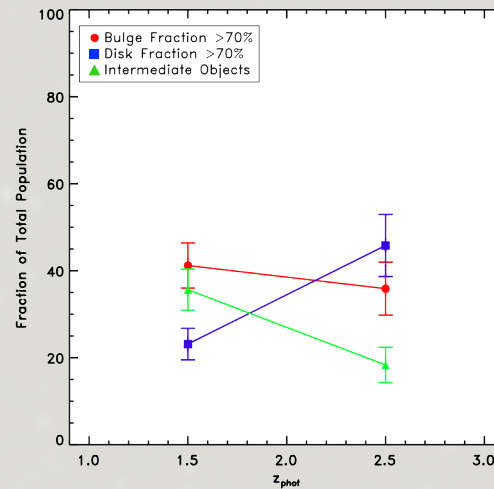
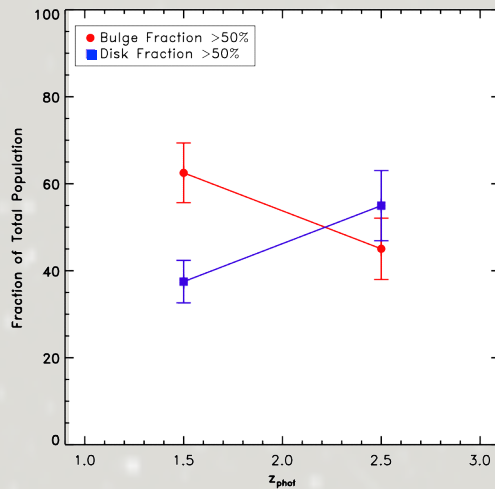
B/T > 0.7  
D/T > 0.7

B/T > 0.9  
D/T > 0.9

fractions  
by light

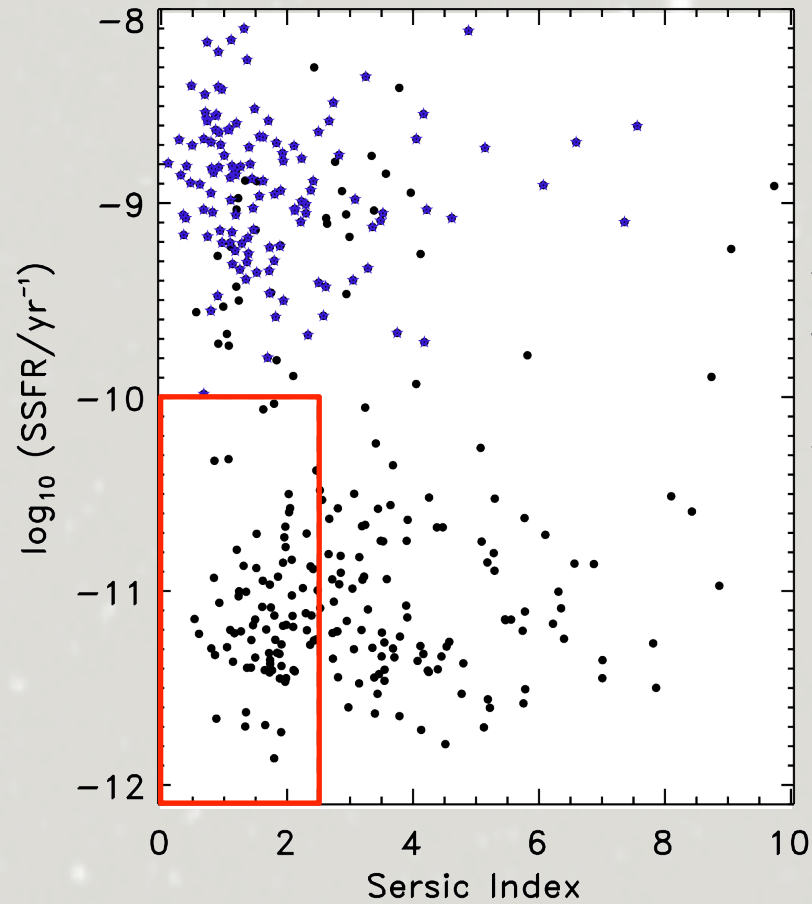


fractions  
by mass

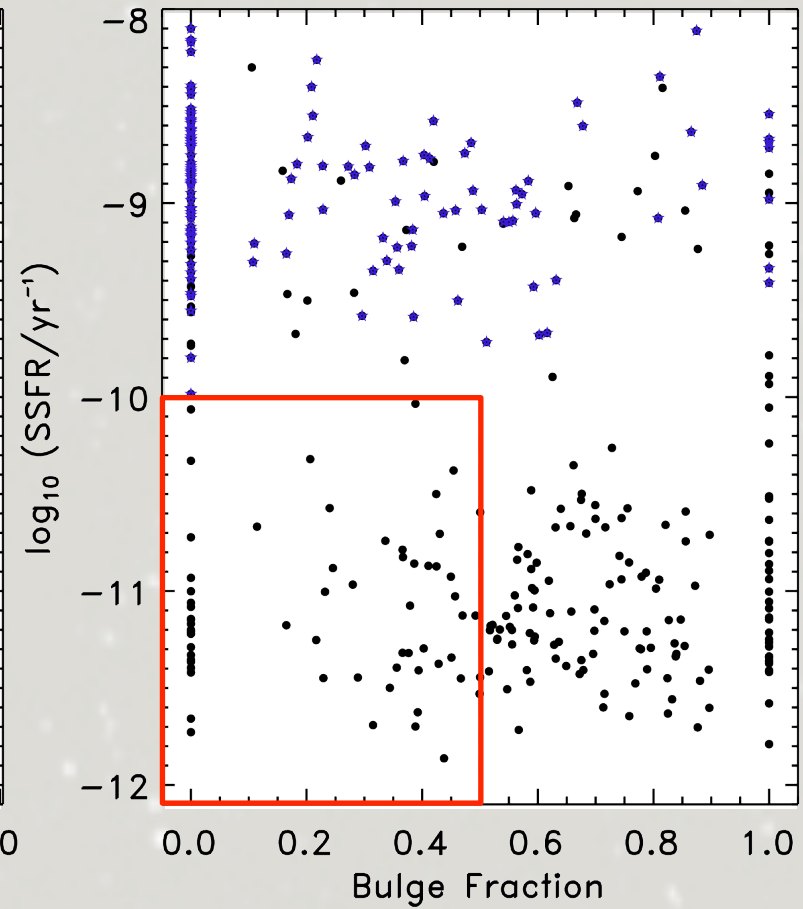


# Morphology and Star-formation Rates

$40 \pm 7\%$



$25 \pm 6\%$

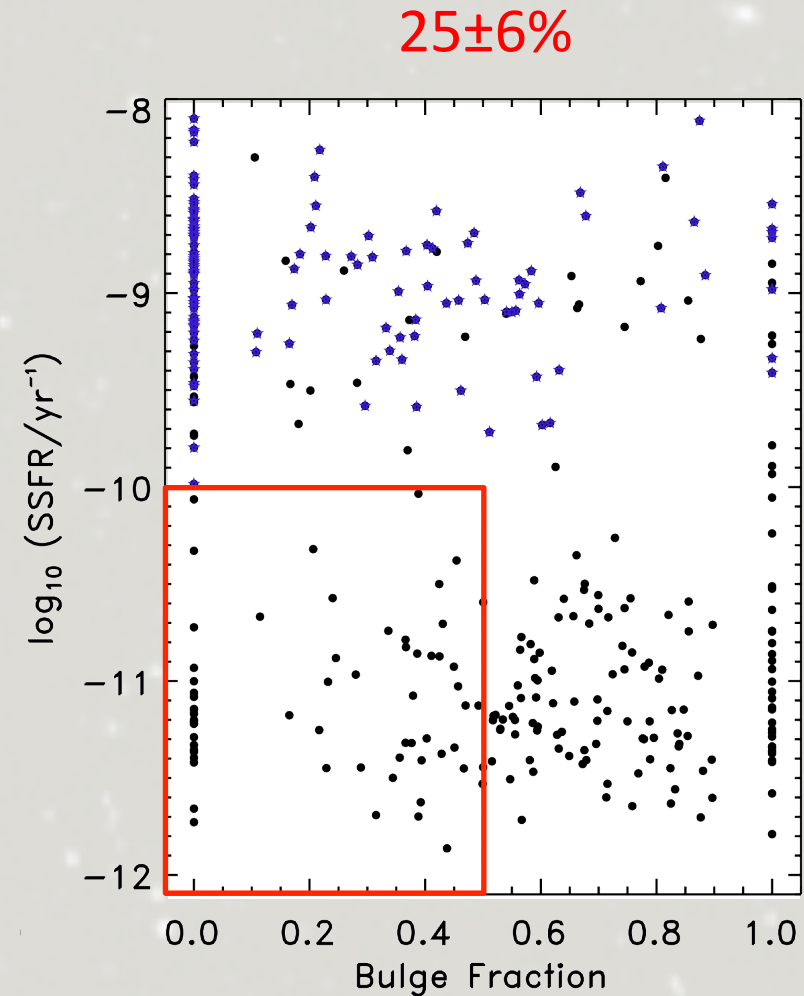




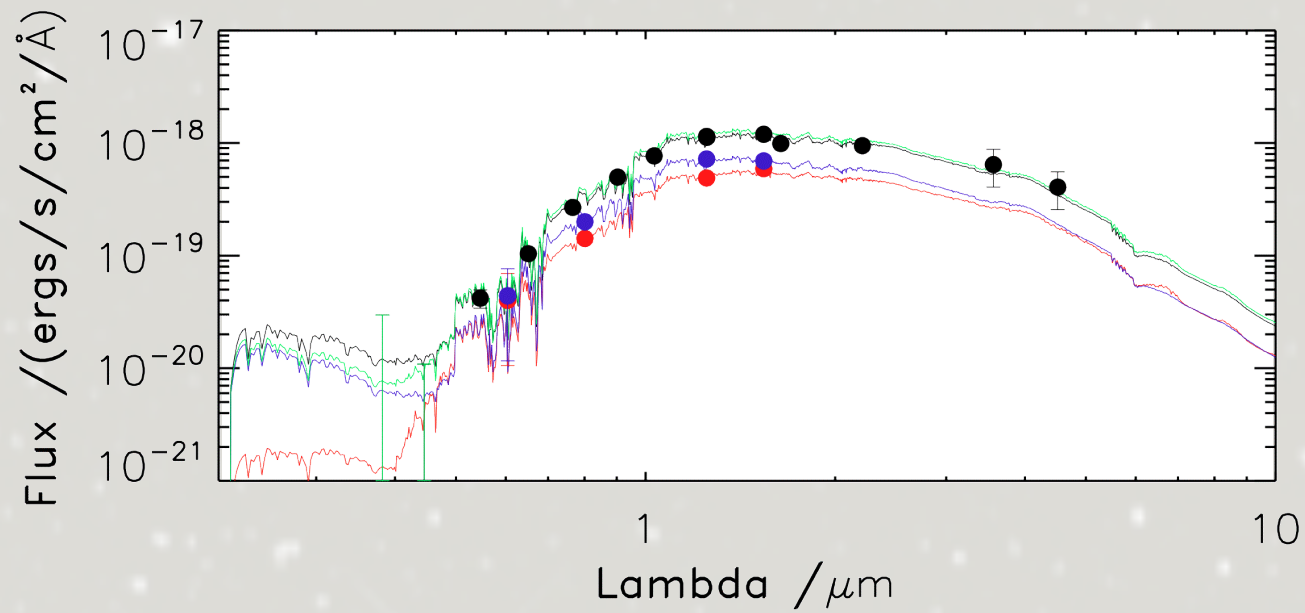
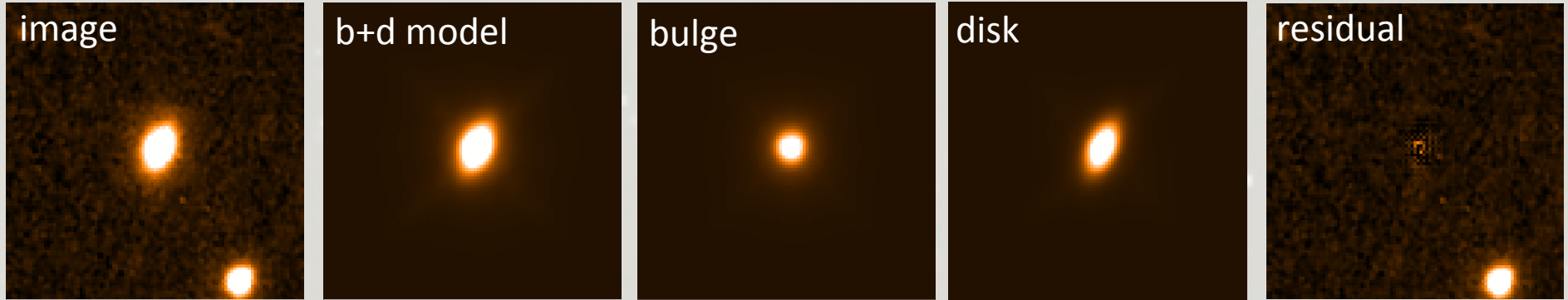
# Morphology and Star-formation Rates

- 61/184 passive galaxies are disk-dominated by bulge light fraction.
- Out of those : 22 are “pure” disks, 39 have best-fit bulge+disk models.
- 30/39 have double component SED fits.
- Only 11/30 are classified as passive disk-dominated galaxies by all criteria.

$18 \pm 4\%$



# Passive Disks



# Conclusions

- Above  $z=2$  massive galaxies become increasingly disk-dominated systems, both in terms of their light and mass fractions.
- Pure bulges are yet to emerge by  $z=1$ .
- $1 < z < 3$  marks the era of the rise of S0 galaxies.
- A significant fraction of passive galaxies are disk-dominated.
- Star-formation quenching is not simply connected to morphological transformations.

## Cold Streams:

Keres et al. 2005  
Dekel et al. 2009a  
Birnboim & Dekel 2003  
Dekel & Birnboim 2006

## Violent Disk Instabilities:

Dekel et al. 2009b  
Ceverino et al. 2010  
Cacciato et al. 2012





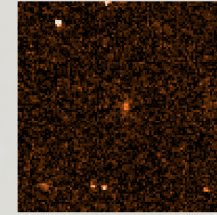
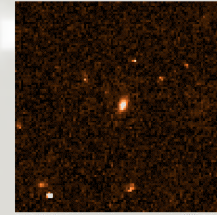
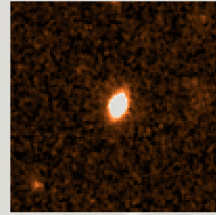
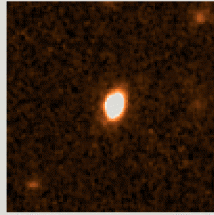
F160W

F125W

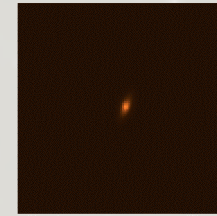
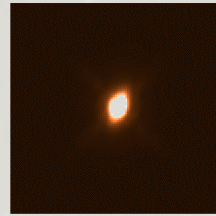
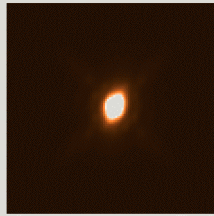
F814W

F606W

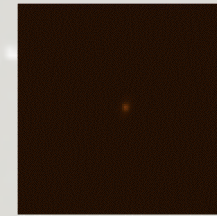
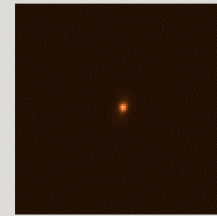
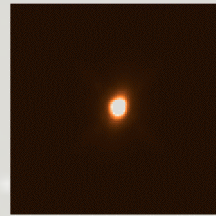
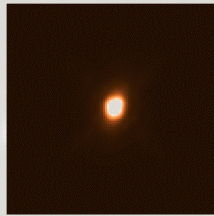
image



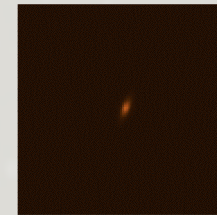
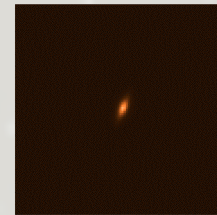
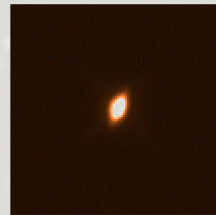
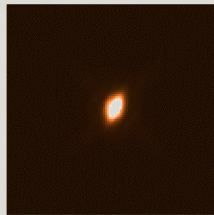
bulge+disk  
model



bulge  
model



disk  
model



residual

