

Gas kinematics in powerful radio galaxies at $z \sim 2$: Energy supply from star formation, AGN, and radio jets^{★,★★}

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ABSTRACT

We compare the kinetic energy and momentum injection rates from intense star formation, bolometric AGN radiation, and radio jets with the kinetic energy and momentum observed in the warm ionized gas in 24 powerful radio galaxies at $z \sim 2$. These galaxies are among our best candidates for being massive galaxies near the end of their active formation period, when intense star formation, quasar activity, and powerful radio jets all co-exist. All galaxies have VLT/SINFONI imaging spectroscopy of the rest-frame optical line emission, showing extended emission-line regions with large velocity offsets (up to 1500 km s^{-1}) and line widths (typically $800\text{--}1000 \text{ km s}^{-1}$) consistent with very turbulent, often outflowing gas. As part of the HeRGÉ sample, they also have FIR estimates of the star formation and quasar activity obtained with *Herschel*/PACS and SPIRE, which enables us to measure the relative energy and momentum release from each of the three main sources of feedback in massive, star-forming AGN host galaxies during their most rapid formation phase. We find that star formation falls short by factors $10\text{--}1000$ of providing the energy and momentum necessary to power the observed gas kinematics. The obscured quasars in the nuclei of these galaxies provide enough energy and momentum in about half of the sample, however, only if both are transferred to the gas relatively efficiently. We compare with theoretical and observational constraints on the efficiency of the energy and momentum transfer from jet and AGN radiation, which favors the radio jets as main drivers of the gas kinematics.

Key words. galaxies: kinematics and dynamics – galaxies: high-redshift – galaxies: active – galaxies: jets – quasars: general – galaxies: star formation

1. Introduction

Powerful radio galaxies (HzRGs) at high redshift ($z \gtrsim 2$) are ideal targets to study the late formation stages of massive galaxies in the early Universe. They have high stellar (e.g., De Breuck et al. 2001, 2010; Seymour et al. 2007) and dynamical masses (Villar-Martín et al. 2003; Nesvadba et al. 2007), and often high star formation rates (SFR) of up to $1000 M_{\odot} \text{ yr}^{-1}$ (Archibald et al. 2001; Reuland et al. 2004; Drouart et al. 2014), with implied formation times of a few 100 Myr. They host luminous, obscured quasars (e.g., Carilli et al. 2002; Overzier et al. 2005; Drouart et al. 2014), and have powerful radio jets (e.g., Carilli et al. 1997; Pentericci et al. 2000), indicating that they are the host galaxies of some of the most powerful active galactic nuclei. Their black hole masses fall near the upper end of the mass function of supermassive black holes in nearby galaxies (Nesvadba et al. 2011a), and scale with the mass of their host galaxies in a fairly similar way to nearby galaxies which fall onto the local black hole–bulge mass relationship, suggesting

they must be near the end of their active formation period. Drouart et al. (2014) argued that the black holes of HzRGs will outgrow the plausible mass range for supermassive black holes even in very massive galaxies if their growth continues for more than a few 10^7 yr, further highlighting that we are observing these sources at a critical moment in their evolution. The same is suggested by their high stellar masses, which exceed the amounts of remaining molecular gas by factors of 10 or more (Seymour et al. 2007; De Breuck et al. 2010; Emonts et al. 2014), limiting their potential future growth in stellar mass.

HzRGs are often surrounded by extended nebulosities of high surface-brightness, warm ionized gas (e.g., Villar-Martín et al. 2003; Nesvadba et al. 2008a) with sizes of up to about 60 kpc, and irregular gas kinematics with velocity offsets and line full widths at half maximum (FWHMs) of up to 1000 km s^{-1} . In very powerful HzRGs, these velocities are above the escape velocity from the gravitational potential of massive galaxies, suggesting that this gas is outflowing (Villar-Martín et al. 2003; Nesvadba et al. 2006, 2008a). In galaxies with more moderate jet power, smaller scale outflows and turbulence seem to be dominant (Collet et al. 2016; Nesvadba et al. 2016, N16 hereafter). The high surface-brightness emission line regions are in most cases elongated along the radio jet axis, and their sizes typically do not exceed the jet size. Dynamical times are comparable to the radio-jet

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lifetime. This supports the interpretation that these are outflows of ambient, warm ionized gas that has been entrained by the expanding “cocoon” of hot, shocked gas inflated by the radio jet (Nesvadba et al. 2006, 2008a).

However, powerful radio galaxies are complex environments where star formation and bright quasar activity co-exist with the powerful radio jets, and it has so far been impossible to compare the possible relative contribution from star formation and quasar activity to the gas kinematics in these systems directly. Star formation as a possible driver of extended, super-galactic bubbles has been suggested by, e.g., McCarthy (1999), Taniguchi et al. (2001), Zirm et al. (2005), and Humphrey et al. (2009). Some of this star formation could be extended in the halo surrounding the radio galaxy itself (e.g., Hatch et al. 2008, 2013).

Estimating star formation rates in high-redshift radio galaxies is observationally challenging. Common star formation tracers like the UV continuum, bright optical emission lines like H α , or rest-frame mid-infrared photometry probing polycyclic aromatic hydrocarbons (PAHs) are generally not reliable tracers of star formation in active galactic nucleus (AGN) host galaxies because they are either contaminated or even dominated by heating through AGN photons and shocks. Although 850 μm photometry with SCUBA has been obtained in the past, and provided interesting qualitative evidence of star formation in HzRGs, deriving robust quantitative constraints has been difficult because of the uncertain spectral shapes and dust temperatures in HzRGs. In addition, the large beam of single-dish measurements raises the dangers of confusing star formation within the radio galaxy with star formation in nearby companion galaxies that fall within the same beam (e.g. De Breuck et al. 2001; Ivison et al. 2008; Nesvadba et al. 2009; Drouart et al. 2014; Emonts et al. 2014). This risk is particularly high in HzRGs since they are predominantly probing dense environments, presumably forming galaxy clusters (Hatch et al. 2009). Furthermore, by analogy with *Spitzer* and *Herschel* observations of nearby radio galaxies, we suspect that even at rest-frame wavelengths of 100–200 μm , as probed in the 850 μm band in galaxies at redshifts $z \sim 2$ –3, the far-infrared (FIR) continuum can still be dominated by the AGN, not the star formation (Tadhunter et al. 2014).

Having a sample of HzRGs with good mid- to far-infrared coverage to disentangle dust heating through AGN radiation and star formation is therefore necessary to study the role of feedback from each process in powerful high-redshift radio galaxies. In addition, we need good constraints on the gas kinematics, ideally over the whole projected surface of the nebulosities.

We present an analysis of the gas energetics in a sample of 24 HzRGs at $z \gtrsim 2$, which have good infrared constraints of obscured AGN and star formation activity, and also rest-frame optical imaging spectroscopy of warm ionized gas. These galaxies are the subset of sources from the HeRGÉ project (Drouart et al. 2014; Seymour et al. 2012), which are also part of our SINFONI study of 49 HzRGs at $z \gtrsim 2$ (Nesvadba et al. 2006, 2007, 2008b, 2011a; Collet et al. 2015, 2016; N16). Our sources have AGN luminosities in the range of $L_{\text{bol,AGN}} = \text{few} \times 10^{13} L_{\odot}$ and radio power of $\geq 10^{27} \text{ W Hz}^{-1}$. They are among the most luminous obscured, radio-loud quasars, and are also among the best studied radio galaxies at their epoch. In our previous work we provided a number of observations suggesting that the jet most likely dominates the kinetic energy injection into the gas, including a close alignment between the jet axis and the kinematic and morphological major axis of the gas, and velocity jumps or elevated line widths near bright radio features in some galaxies. However, as encouraging as these findings are, it is not sufficient to show that the jet can plausibly accelerate the gas; we also need to show

that alternative mechanisms are less likely. This has now become possible thanks to the HeRGÉ constraints, and is the goal of the present study.

Seymour et al. (2012) have already carried out a similar analysis for one galaxy of our set. For the Spiderweb galaxy MRC 1138–262 at $z = 2.16$, they estimated a bolometric power of the AGN of $L_{\text{bol,AGN}} \sim 7 \times 10^{13} L_{\odot}$ and argued that this would suffice to power the outflow with kinetic energy of $L_{\text{kin}} \sim 8 \times 10^{12} L_{\odot}$ observed in MRC 1138–262 by Nesvadba et al. (2006). Hence the AGN and the radio jet would seem energetically capable of driving the gas kinematics in this galaxy, if about 10% of the bolometric luminosity of the AGN was transformed into kinetic energy of the gas. In turn, the kinetic energy released by the star formation in MRC 1138–262, although one of the most vigorously star-forming high-redshift radio galaxies known, would not be sufficient (Nesvadba et al. 2006; Seymour et al. 2012).

Building upon these results for a first galaxy, we present a direct comparison of the energy injection rates from star formation, quasar bolometric power and radio jets. We expand the previous results by including all 24 galaxies with *Herschel* photometry and VLT/SINFONI imaging spectroscopy, and by analyzing the energy as well as the momentum injection into the gas. We discuss our results in light of recent feedback models.

Throughout the paper we adopt a flat $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$ concordance cosmology with $\Omega_{\text{M}} = 0.3$ and $\Omega_{\Lambda} = 0.7$.

2. Mid- to far-infrared SEDs and IR luminosities

The set of galaxies we study here is a subset of the HeRGÉ set of 70 distant radio galaxies at redshifts $1 < z < 5.2$ with *Herschel* PACS and SPIRE photometry in the far-infrared (Drouart et al. 2014; Seymour et al. 2012). These galaxies were also observed with SINFONI and are therefore at redshifts where bright emission lines fall into the near-infrared atmospheric windows (predominantly $z \sim 2.0$ –2.6, and $z \sim 3.0$ –3.6 N16), and at either southern or equatorial declinations. All the galaxies also have *Spitzer* mid-infrared photometry (Seymour et al. 2007; De Breuck et al. 2010), and sub-mm photometry obtained with either JCMT/SCUBA (Archibald et al. 2001; Reuland et al. 2004) at 850 μm , or APEX/LABOCA (Drouart et al. 2014) at 870 μm . In total, their infrared to submillimeter spectral energy distributions cover the wavelength range between 16 and 870 μm .

Drouart et al. (2014) used DecompIR (Mullaney et al. 2011) to decompose these SEDs into an AGN and a starburst component, using the empirical templates of Mullaney et al. (2011) for both components. Mullaney et al. (2011) provide five different empirical templates for the starburst component. Most of the galaxies in our present sample are best fit by their SB 2 template, i.e., the infrared spectral energy distribution of NGC 7252 (Drouart et al. 2014, their Table 6). The 24 sources we are concerned with here have infrared luminosities from AGN and star formation between 1.6 and $24.5 \times 10^{12} L_{\odot}$ and 2.2 and $10.5 \times 10^{12} L_{\odot}$, respectively (Table 6 of Drouart et al. 2014), integrating over wavelengths 8–1000 μm in the rest-frame. Spectral decompositions for individual sources are shown in Fig. D.1 of Drouart et al. (2014).

3. Kinetic energy and momentum of the warm ionized gas

All the galaxies have also been observed with the SINFONI imaging spectrograph on the Very Large Telescope of ESO

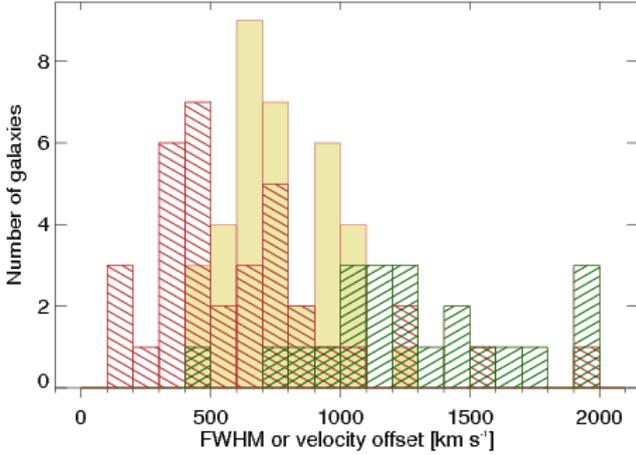


Fig. 1. Average (yellow shaded histogram) and maximum (dark green hatched histogram) FWHMs of our sources, and of their velocity offsets (dark red hatched histogram).

(Nesvadba et al. 2006, 2008a; Collet et al. 2016, N16). SINFONI is an image slicer with $8'' \times 8''$ field of view operating in the near-infrared J , H , and K bands with spectral resolving power between $R = 1500$ and 3000 . We mainly used the seeing limited mode with spatial sampling of $125 \text{ mas} \times 250 \text{ mas}$ to observe the bright rest-frame optical emission lines, in particular $[\text{OIII}]\lambda\lambda 4959, 5007$ and $\text{H}\alpha$, depending on redshift.

Figure 1 summarizes the kinematic properties of the warm ionized gas in our galaxies, which has already been presented in greater detail by N16, in the form of three histograms: we show the average and maximal FWHM measured from $[\text{OIII}]\lambda 5007$ in our SINFONI cubes for each galaxy, and the maximum velocity offset in the extended emission line gas. The average FWHM ranges from $FWHM_{\text{avg}} = 400 \text{ km s}^{-1}$ to 1100 km s^{-1} , and reaches 2500 km s^{-1} in USS 1243+036. It represents the average of all local FWHMs measured typically from $[\text{OIII}]\lambda 5007$ in all spatial pixels in a given galaxy where the signal-to-noise ratio exceeds three. The maximum FWHM measured in small regions of our sources are between 700 km s^{-1} and 1950 km s^{-1} , and are typically about 500 km s^{-1} greater than the average FWHM (N16). The velocity offsets are between 100 km s^{-1} and 1300 km s^{-1} for most sources. We note that even if we corrected the velocity offsets by fiducial factors 2–3 to account for inclination effects, the radial velocities of the gas would typically not exceed the FWHM line widths.

A potential caveat of these estimates might be that $[\text{OIII}]\lambda 5007$ is probing a highly ionized state of a trace element, and may therefore not be a representative tracer of the dominant mass component. However, N16 compared the velocity fields of $[\text{OIII}]\lambda 5007$ and $\text{H}\alpha$ in the nine galaxies where both are emitted from spatially well-extended regions and at good signal-to-noise ratios, finding little cause for alarm, at least at the spatial resolutions of about 5–10 kpc reached with this study.

Our imaging spectroscopy data provide several constraints that we can use to estimate the kinetic energy in the warm ionized gas of our sources. We follow Nesvadba et al. (2006) in estimating the energy necessary to inflate a hot bubble adiabatically to the sizes, r , and velocities, Δv , as observed:

$$dE/dt = 1.5 \times 10^{46} \times r^2 \times \Delta v^3 \times n_0 \text{ erg s}^{-1} \quad (1)$$

where r is the radius of the bubble in units of 10 kpc. Δv is given in units of 1000 km s^{-1} , and we adopt $n_0 = 1 \text{ cm}^{-3}$ for the density of the gas into which the bubble expands.

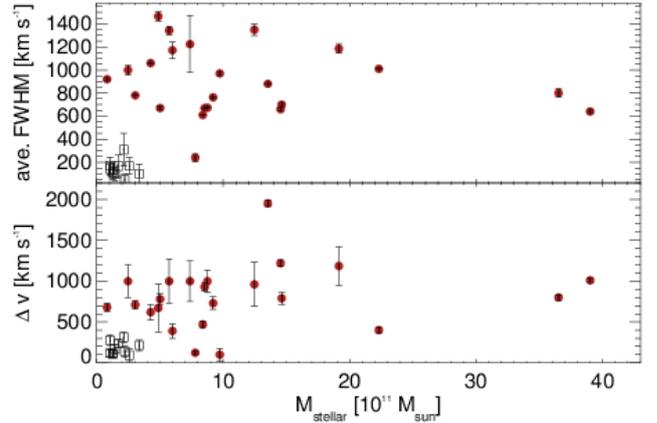


Fig. 2. Emission-line FWHM (*top*) and velocity offset (*bottom*) as a function of stellar mass, as estimated by Seymour et al. (2007) and De Breuck et al. (2010) from *Spitzer* near- to mid-infrared photometry. The red dots show the galaxies from our sample, the small black empty squares the stellar-mass selected sample of Buitrago et al. (2014) without powerful AGN for comparison. We do not find a correlation for our sample in either relationship.

This estimate does not depend on the total gas mass, and is therefore not affected by gas in other phases than the warm ionized phase, which may contribute significantly to the total mass in the outflow (e.g., Cicone et al. 2014). The only requirement is that the gas kinematics are representative of the terminal outflow velocity, which should be the case, because they subtend large volumes surrounding the galaxy relatively evenly and have moderate densities compared to the colder, and perhaps dominant, gas phases (N16). Table 1 shows that the resulting kinetic energy in our targets is between $\log E_{\text{gas}} = 43.0$ and 47.5 erg s^{-1} .

4. Stellar masses and gravitational motion

Before turning to the potential astrophysical mechanisms that may power the gas kinematics in our HzRGs, we first present our arguments for why gravitational motion is unlikely to play a large role. Seymour et al. (2007) and De Breuck et al. (2010) used *Spitzer* and ground-based photometry to decompose the optical to infrared spectral energy distributions of our sources into their stellar and AGN dust components, obtaining robust stellar mass estimates (or at least upper limits) for most of them. Most HzRGs fall into a small range of (high) stellar masses around $\log M_{\text{stellar}}/M_{\odot} = 11.5 \pm 0.3$.

In Fig. 2, we show the total velocity offsets and average FWHMs of the gas as a function of the stellar mass derived by Seymour et al. (2007) and De Breuck et al. (2010) for the sources we study here. We find no evidence of a relationship between stellar mass and velocity or FWHM, as measured across the extended emission-line regions, and consider this additional evidence that gravity does not play a large role in the gas motion in our sources. We also show the sources of Buitrago et al. (2014), to our knowledge the only available stellar mass selected sample of galaxies in the early Universe with rest-frame optical imaging field spectroscopy and with $M_{\text{stellar}} > 10^{11} M_{\odot}$. This adds to the arguments already presented by Collet et al. (2016) and N16, namely that the velocity and dispersion maps of the gas are very complex and inconsistent with simple rotating disks, that line widths are systematically and significantly greater than in quiescent high- z bulges with similar stellar mass, and that their ratios of velocity gradients to line widths are inconsistent with

Table 1. Coordinates, redshifts, gas kinetic energy, AGN bolometric power, and kinetic energy from star formation and radio jet for the sources in this sample.

Source	RA (J2000)	Dec (J2000)	z	dE/dt_{gas} [erg s ⁻¹]	$dE/dt_{\text{AGN,bol}}$ [erg s ⁻¹]	dE/dt_{SF} [erg s ⁻¹]	dE/dt_{jet} [erg s ⁻¹]
MRC 0114–211	01:16:51.4	-20:52:07	1.41	43.9 ± 0.3	46.6 ± 0.1	43.4 ± 0.1	47.6 ± 0.3
TNJ 0121+1320	01:21:42.7	+13:20:58	3.52	44.5 ± 0.3	<46.9	44.0 ± 0.1	47.0 ± 0.3
MRC 0156–252	01:58:33.6	-24:59:31	2.02	46.6 ± 0.3	47.3 ± 0.1	<43.6	47.0 ± 0.3
TNJ 0205–2242	02:05:10.7	+22:42:50	3.51	44.5 ± 0.3	<46.8	<43.7	47.0 ± 0.3
MRC 0211–122	02:14:17.4	-11:58:46	2.34	43.0 ± 0.3	47.3 ± 0.1	43.7 ± 0.2	47.0 ± 0.3
MG 0251–273	02:53:16.7	-27:09:13	3.16	45.2 ± 0.3	46.8 ± 0.1	<43.9	47.2 ± 0.3
MRC 0316–257	03:18:12.1	-25:35:10	3.14	45.7 ± 0.3	<46.6	44.0 ± 0.1	47.4 ± 0.3
MRC 0406–244	04:08:51.4	-24:18:17	2.44	46.1 ± 0.3	47.1 ± 0.1	44.0 ± 0.1	47.5 ± 0.3
PKS 0529–549	05:30:25.4	-54:54:22	2.58	44.5 ± 0.3	46.8 ± 0.1	44.1 ± 0.1	47.5 ± 0.3
5C 7.269	08:28:38.8	+25:28:27	2.22	44.0 ± 0.3	<46.9	<43.5	46.6 ± 0.3
TXS 0828+193	08:30:53.4	+19:13:16	2.57	46.0 ± 0.3	47.6 ± 0.1	43.5 ± 0.2	46.9 ± 0.3
3C257	11:23:09.4	+05:30:18	2.46	45.0 ± 0.3	47.2 ± 0.1	43.7 ± 0.1	47.8 ± 0.3
MRC 1138–262	11:40:48.3	-26:29:10	2.16	46.5 ± 0.3	47.4 ± 0.1	43.7 ± 0.1	47.4 ± 0.3
USS 1243+036	12:45:38.4	+03:23:21	3.57	46.7 ± 0.3	47.3 ± 0.1	43.8 ± 0.2	47.5 ± 0.3
USS 1410–001	14:13:15.1	-00:23:00	2.36	45.7 ± 0.3	47.1 ± 0.1	<44.10.1	47.0 ± 0.3
MRC 1558–003	16:01:17.3	-00:28:48	2.53	45.7 ± 0.3	47.3 ± 0.1	<43.5	47.3 ± 0.3
USS 1707+105	17:01:06.5	+10:31:06	2.35	45.4 ± 0.3	46.5 ± 0.1	<43.7	47.2 ± 0.3
TNJ 2007–1316	20:07:53.2	-13:16:44	3.84	44.2 ± 0.3	47.3 ± 0.1	43.9 ± 0.1	47.4 ± 0.3
MRC 2025–218	20:27:59.5	-21:40:57	2.63	44.8 ± 0.3	<46.6	<43.9	47.2 ± 0.3
MRC2104–242	21:06:58.2	-24:05:11	2.49	46.2 ± 0.3	46.8 ± 0.2	43.8 ± 0.1	47.3 ± 0.3
4C23.56	21:07:14.8	+23:31:45	2.48	46.1 ± 0.3	47.8 ± 0.1	<43.8	47.3 ± 0.3
MRC 2144+1928	21:44:07.5	+19:29:15	3.59	45.7 ± 0.3	47.3 ± 0.1	43.7 ± 0.1	47.5 ± 0.3
MRC2224–273	22:27:43.2	-27:05:02	2.15	43.4 ± 0.3	46.7 ± 0.1	43.4 ± 0.1	46.7 ± 0.3

Notes. FIR luminosities from AGN and star formation have been given by [Drouart et al. \(2014\)](#), and 1.4 GHz radio power in [N16](#).

the ellipticities for both fast and slow rotators in the *Atlas*^{3D} sample of nearby galaxies.

The broad line widths and high velocities further suggest that an additional energy and momentum input mechanism must be present in the HzRGs. In the following sections we use our observational constraints on star formation rates, AGN bolometric power, and radio jets to further discuss which of these three dominates.

5. Power of star formation

Intense star formation in galaxies releases kinetic energy and momentum into the ambient gas in form of stellar winds from young stars and through expanding shells of supernova remnants, which interact, mix, and thermalize, producing smooth, galaxy-wide outflows (e.g., [Heckman et al. 1990](#); [Lehnert & Heckman 1996a; 1996b](#)). Observations and simulations have now reached a broad consensus that about 40% of the initial 10^{51} erg of kinetic energy released by a supernova is carried by galactic winds (e.g., [Dalla Vecchia & Schaye 2008](#); [Sharma et al. 2014](#); [Veilleux et al. 2005](#); [Strickland & Heckman 2009](#)), thus releasing about $1\text{--}2 \times 10^{49}$ erg of kinetic energy per solar mass of stars formed into the interstellar medium ([Dalla Vecchia & Schaye 2008](#)). The precise estimate depends most strongly on the assumptions about the initial mass function of star formation and the details of the thermalization and mass loading process. Here we adopt an energy release of 10^{49} erg per solar mass of star formation.

We estimate the kinetic energy released by star formation in our HzRGs from the starburst component of the far-infrared luminosity measured by [Drouart et al. \(2014\)](#), and calculate a star

formation rate by setting $SFR = 4.5 \times 10^{-44} L_{\text{FIR}}$ ([Kennicutt 1998](#)). The SFR is given in $M_{\odot} \text{yr}^{-1}$, and the FIR luminosity in erg s^{-1} . [Kennicutt \(1998\)](#) derived this relationship for a Salpeter initial mass function; we divide the star formation rates by a factor of 1.8 to match the Chabrier initial mass function more commonly used today ([Chabrier 2003](#)).

With star formation rates between 211 and $550 M_{\odot} \text{yr}^{-1}$, and using the above mentioned scaling that supernovae release on average 1×10^{49} erg of kinetic energy per solar mass of stellar mass formed, we find kinetic energies between 1.7×10^{44} erg s^{-1} and 2.2×10^{45} erg s^{-1} . This corresponds to $0.4\text{--}5.8 \times 10^{11} L_{\odot}$.

To determine the momentum produced by the star formation in our sources, we rely on the recent empirical study of powerful starburst-driven winds in low-redshift galaxies by [Heckman et al. \(2015\)](#), who find that the momentum injection per solar mass of new stars formed is about 4.8×10^{33} dyn. This includes the contribution of ram pressure from the hot wind medium as well as that of radiation pressure. They also find that the starburst must inject about $10\times$ more momentum into the gas than is necessary to unbind the wind, in order to produce clear outflow signatures. In our targets this corresponds to 9×10^{35} dyn to 6×10^{36} dyn. Values for individual galaxies are listed in [Table 1](#).

[Figure 3](#) shows the kinetic energy and momentum of the gas as a function of the energy and momentum that are released by star formation with the above estimates. Nearly all sources are above the solid black line for at least one quantity, with only one exception, USS 0211–122 at $z = 2.1$. This galaxy has unusually narrow emission lines for our overall sample, and a distinctive unresolved broad component, which is also unusual for our sample ([N16](#)). [Vernet et al. \(2001\)](#) also found it was an outlier

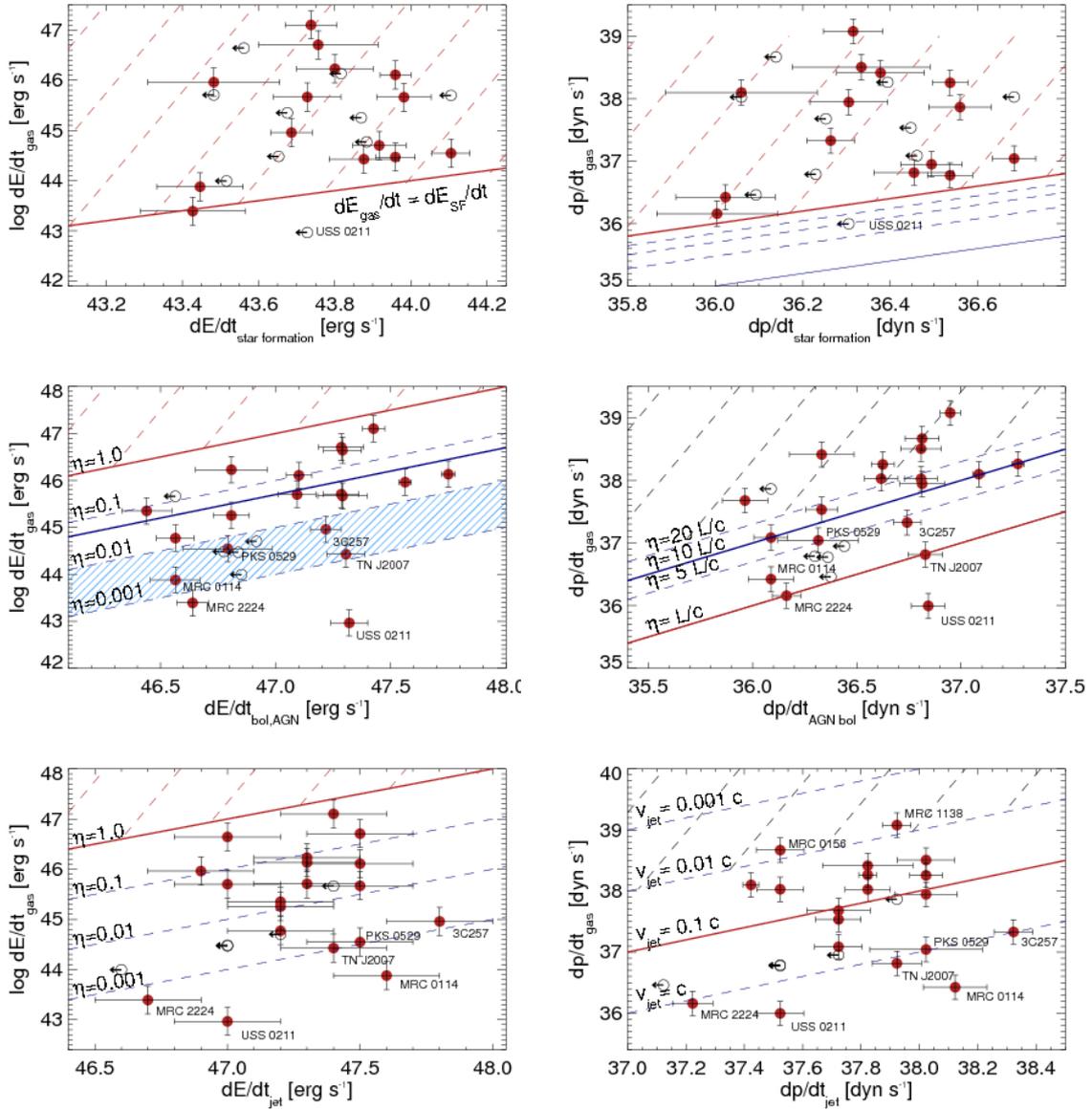


Fig. 3. Relationships between the observed kinetic energy (*left*) and momentum (*right*) of the gas, for star formation (*top*), AGN radiation (*center*) and radio jets (*bottom*). Filled red and empty black dots show galaxies with FIR detections and upper limits, respectively. Error bars show the measurement uncertainties. Red wide-spaced dashed regions are excluded by energy and momentum conservation, black wide-spaced dashed regions by empirical constraints on the maximum momentum transfer from AGN radiation and radio jet, respectively. In the *top right panel*, the red and blue solid lines show momentum transfer of 100 and 10% of that provided by star formation, respectively. Heckman et al. (2015) argued that 10% efficiency is most likely. Thin dashed dark blue lines show transfer rates of 30, 50, and 70%. In the *middle left panel*, the red line shows an energy transfer of 100%, the blue line 5% required by cosmological models, the hatched region of 0.1–1% of L_{bol} to gas kinetic energy. The latter is favored by observations of FeLoBALs (Moe et al. 2009; Dunn et al. 2010; Bautista et al. 2010) and hydrodynamic models taking radiative transfer explicitly into account (Bieri et al. 2017; Krumholz & Thompson 2012). The same observations and models also favor an average momentum transfer of L/c (red line in the *middle right panel*), whereas more optimistic models find $L/c \sim 10$ (dark blue solid line; Zubovas & King 2012; Faucher-Giguère & Quataert 2012). In the *bottom left panel*, the red solid line represents an energy transfer of 100% from jet to gas, and in the *bottom right panel* the momentum transfer for the most likely jet expansion velocity of 0.1c. Dashed dark blue lines show the momentum input from the jet for different assumptions of the jet advance speed, assuming that no momentum is lost to other effects. It is therefore not impossible to find galaxies below the line for a jet expansion at the speed of light, c (lowest dark-blue dashed line), due to geometry, gas clumpiness, and other effects.

compared to other HzRGs, with unusually high levels of polarization consistent with low levels of star formation, high nitrogen abundance suggesting high levels of secondary nitrogen production, and perhaps a very advanced evolutionary state.

The remaining 23 galaxies require more energy and momentum input than can be provided by star formation, by up to three orders of magnitude, even if this energy and momentum were

transferred into the interstellar medium at 100% efficiency. This is likely not a realistic assumption. Observations of starburst-driven winds suggest that efficiencies are more in the range of a few times 10% (Strickland & Heckman 2009) for the energy, and about 10% for the momentum transfer (Heckman et al. 2015), larger than the potential systematic uncertainties. The latter efficiency is shown as a dashed black line in the top right panel

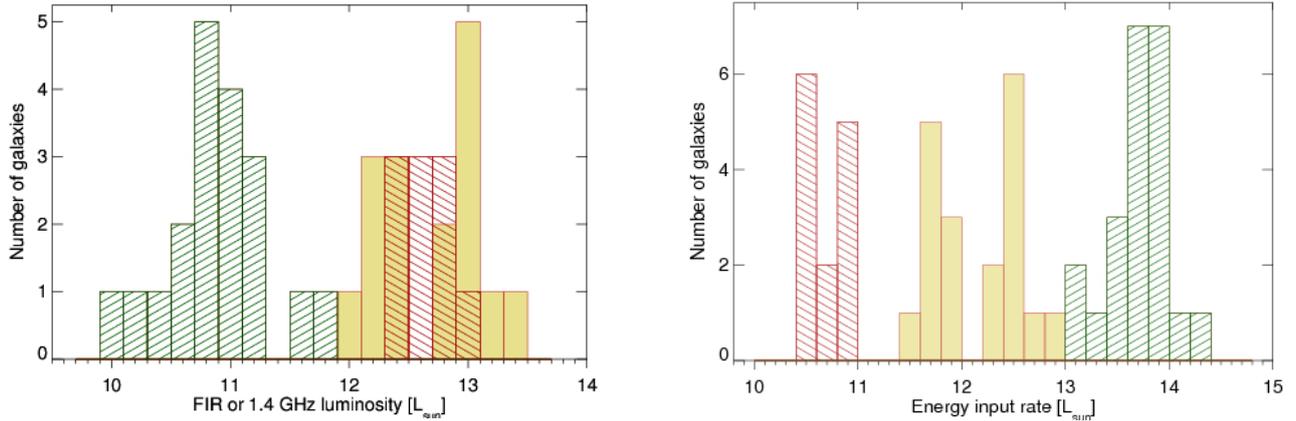


Fig. 4. *Left:* FIR luminosity between 8 and 1000 μm , and the 1.4 GHz monochromatic power, i.e., the quantities used to estimate the kinetic energy of the different possible drivers discussed in this paper. The yellow filled, red hatched, and green hatched diagrams correspond to bolometric AGN power, star formation, and radio power, respectively. We only include sources with *Herschel* detections. *Right:* same histograms showing the kinetic energy corresponding to the FIR luminosity and monochromatic radio power at 1.4 GHz, derived with the methods described in Sects. 5 and 6.

of Fig. 3. Many theoretical studies suggest values that are even lower (e.g., Krumholz & Thompson 2012; Bieri et al. 2017). As an additional caveat, recent ALMA arcsec-resolution observations have shown that most of the FIR emission from star formation does not necessarily come from the radio galaxy itself, but from nearby companion galaxies (Gullberg et al. 2016, see also Ivison et al. 2008; N16). This would also suggest that we might overestimate the energy and momentum input rates shown in Fig. 3. Taking these arguments into account only reinforces our finding that star formation in these galaxies does not produce sufficient kinetic energy or momentum to power the gas kinematics as observed.

6. Bolometric AGN radiation

Observations of broad quasar absorption lines have provided secure evidence of fast outflows within the nuclei of powerful AGN for several decades, but observational evidence that such outflows may also extend into the kpc range away from galactic nuclei is being found only now (e.g., Liu et al. 2013a,b; Sun et al. 2014), although observations of broad, blueshifted emission-line wings from the narrow line region have been known for much longer (e.g., Heckman et al. 1984; Greene & Ho 2005; Woo et al. 2016). The bolometric luminosity emitted by powerful quasars during their lifetimes corresponds approximately to the binding energy of their massive host galaxy (e.g., Silk & Rees 1998). Quasars might therefore in principle drive large-scale outflows if the radiated energy and momentum are deposited in the ambient gas at high enough efficiency. Cosmological models of galaxy evolution suggest that about 0.2–0.5% of the rest-mass energy equivalent of a supermassive black hole must be injected to reproduce the observed black-hole bulge (mass) relationships (e.g., Di Matteo et al. 2005), corresponding to about 2–5% of the bolometric luminosity of the AGN for a radiative accretion efficiency onto the black hole of 0.1.

To estimate the bolometric quasar power, we rely on the AGN component of the FIR spectral energy distributions of Drouart et al. (2014, their Table 6), and multiply by a factor of 6 (again following Drouart et al. 2014), to find the bolometric quasar luminosities listed in Table 1 of between 6.6×10^{45} and 5.6×10^{47} erg s^{-1} . These estimates are uncertain by factors of a few (Drouart et al. 2014) for two main reasons: First, the AGN torus model adopted by Mullaney et al. (2011) assumes

relatively high dust temperatures, which would maximize the FIR AGN luminosity by a factor of ~ 2 at the expense of the FIR luminosity from star formation, which would be underestimated by a similar factor. However, if the torus component is not sufficiently extended, or if a significant part of the star formation occurs in companion galaxies, we would overestimate the starburst and underestimate the AGN component (Drouart et al. 2016). Second, as already pointed out by Drouart et al. (2014, their Appendix C), the adopted AGN spectral shape used for the bolometric correction is also uncertain by factors of a few. We adopted their favored bolometric correction factor 6 obtained from the classical QSO template by Elvis et al. (1994); the template of Richards et al. (2006) would have resulted in a bolometric correction factor of 5. For four galaxies from the present sample Nesvadba et al. (2011a) estimated bolometric luminosities from the luminosity and width of the broad nuclear $\text{H}\alpha$ line, finding values between $\log L_{\text{bol}} = 46.6$ and 46.8 erg s^{-1} , consistent with X-ray constraints in the two cases where X-ray constraints were available. Generally, these estimates are in the lower range of the $\log L_{\text{bol}} = 46.6$ to 47.4 erg s^{-1} found with the approach of Drouart et al. (2014). Therefore, we cannot exclude that we overestimate our bolometric AGN luminosities by factors of a few. In the present case, however, we wish to distinguish between the impact of quasar radiation and radio jets, which means that this is a conservative choice that maximizes the potential impact of the quasar radiation.

In Fig. 4 we compare the far-infrared luminosities from the quasar and from star formation in our targets, finding that both are quite similar, about few 10^{12} L_{\odot} for most targets. It is, however, also interesting to compare the kinetic power released by the starburst, estimated in the way described in Sect. 5, with that from the AGN, assuming that 5% of the bolometric AGN power is turned into mechanical energy of the gas. Figure 4 shows that the kinetic energy from the AGN radiation exceeds that produced by star formation by one to two orders of magnitude.

The AGN could therefore unbind the gas at least in some of the sample. This is also shown in the middle left panel of Fig. 3 where we show the kinetic energy of the gas in each AGN host galaxy as a function of the bolometric AGN luminosity. The red solid line shows where the ratio of both is unity, and the dark blue dashed lines indicate conversion efficiencies between 0.1 and 10% between bolometric luminosity and gas kinetic energy. The dark blue line shows the 5% efficiency suggested by the

black-hole bulge mass relationship. If this efficiency is approximately correct, then the AGN should provide sufficient energy to unbind the gas in about two-thirds of our sample.

It is, however, still a matter of active debate among observers and modelers whether feedback efficiencies as high as a few percent are not unrealistically high. Observations of molecular gas in very dense, vigorously star-forming circumnuclear environments of nearby AGN/starburst composites suggest that similarly high conversions are necessary if the AGN is to dominate the observed gas kinematics. In turn, observations of UV absorption lines in FeLoBALs only find efficiencies between 0.1 and 1% (Moe et al. 2009; Dunn et al. 2010; Bautista et al. 2010). FeLoBALs are thought to trace winds with particularly high dust column densities where the energy deposition efficiencies should thus be high over comparably large radii. In Fig. 3 this range corresponds to the light blue dashed region. We label all galaxies that fall into or below this region in this diagram, and we return to a more detailed discussion of their properties below (Sect. 8).

In addition to the energy transfer, the quasar radiation also injects momentum into the gas. Each time a photon is scattered on a gas or dust particle, a momentum of $h\nu/c$ is transferred, leading to a total momentum transfer of $\xi L_{\text{bol}}/c$, where the “momentum boost” ξ corresponds to the number of scatterings suffered by a photon before it escapes, depending on the wavelength of the photon and the optical depth along the path.

In the middle right panel of Fig. 3 we show the momentum of the gas as a function of the momentum from the quasar radiation for different assumptions of ξ between $\xi = 1$ and $\xi = 20$. Most smoothed particle hydrodynamic models find that momentum transfer of up to $20 L_{\text{bol}}/c$ is necessary to explain efficient AGN feedback as implied by the black hole–bulge scaling relations, but make simplified assumptions, for example on the radiative transfer or the energy and momentum transfer from the hot AGN wind onto the ambient gas. The recent models of Zubovas & King (2012) and Faucher-Giguère & Quataert (2012), which combine the effects of the acceleration of gas in an expanding blastwave and radiation pressure, suggest a momentum boost of about $\xi = 10$, highlighted as a blue solid line in Fig. 3.

Analytical work and hydrodynamic models taking radiative transfer explicitly into account find, however, that most AGN host galaxies, except perhaps in the densest star-forming regions in ULIRGs, globally may only reach a much lower total momentum transfer of about L_{bol}/c (red line in Fig. 3; e.g., Murray et al. 2005; Novak et al. 2012; Krumholz & Thompson 2013). In an adaptive-mesh simulation with a simplified but nonetheless explicit radiation transfer method, Bieri et al. (2017) recently found a rapid decrease from about 30 to less than $1 \times L_{\text{bol}}/c$ within the first 10^7 yr after the ignition of the radiative AGN. The reason is that the radiation escapes relatively easily after the AGN has blown low-density channels into the dense gas clouds by which it is initially enshrouded. This is also more in line with the observations of FeLoBALs. Likewise, Veilleux et al. (2013) suggested that molecular outflows from the highest density inner regions of nearby ULIRGs harboring powerful, dust-enshrouded AGN, which are likely driven mainly by radiation pressure, may subside once the AGN has cleared a path through the ambient dust and gas clouds. Nonetheless, radiation pressure and disk winds could perhaps play a role during early feedback phases in our galaxies before the winds have broken away from the direct surroundings of the AGN in the central regions of the host galaxy. While a momentum boost of $\xi = 10$ – 20 would appear sufficient to power the observed outflows in about half of our sample; with $\xi = 1$, the obscured quasars in our

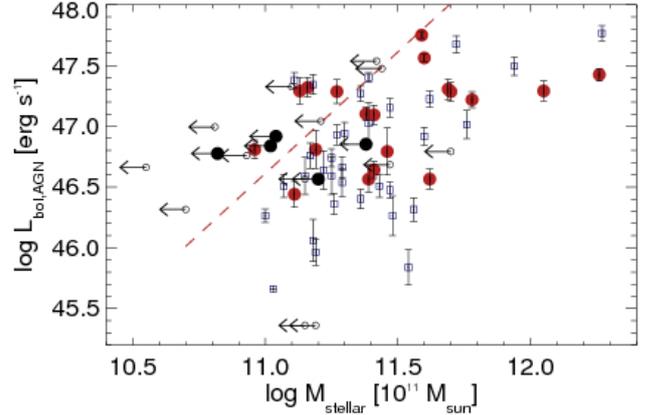


Fig. 5. Critical AGN luminosity for radiatively driven quasar winds in the model of Murray et al. (2005) as a function of stellar mass. Large filled red dots and filled black dots show galaxies from our sample with measurements and upper limit of stellar mass, respectively. Small empty blue boxes and small empty black dots show the HeRGE galaxies without SINFONI observations and with measurements and upper limits on stellar mass, respectively. Quasar luminosities were taken from Drouart et al. (2014) and correspond to their AGN component of the FIR luminosity, which was scaled by a factor of 6 to approximate the bolometric QSO luminosity. The dashed red line shows the critical luminosity at which the momentum carried by the quasar radiation may balance gravity (see Murray et al. 2005, for details).

galaxies would be too weak for all but three sources with measurements of L_{bol} (and one with an upper limit) to explain the observed gas kinematics. We have labeled these galaxies individually in Fig. 3, and discuss their specific properties in more detail in Sect. 8 after examining the role of the radio jets.

The same is found when comparing the AGN bolometric luminosities in our galaxies with those required to balance the gravitational potential of the host galaxy through radiation pressure (but not necessarily, to remove the gas through an outflow). Murray et al. (2005) introduced such a quantity for a model of the combined feedback from the radiation produced by AGN and star formation. They find that in the most fortuitous case, when the surroundings of the AGN are optically thick, the AGN must emit a critical luminosity $L_c = 4f_g c \sigma^4 / G$ that depends on the gas fraction, f_g , and stellar velocity dispersion, σ ; G is the gravitational constant.

Using the stellar mass estimates of Seymour et al. (2007) and De Breuck et al. (2010) typically of a few $10^{11} M_\odot$, and the molecular, atomic, and warm ionized gas mass estimates of, e.g., Emonts et al. (2014), Gullberg et al. (2016), and Nesvadba et al. (2016), we find that the typical gas fractions in HzRGs are about 10%. Adopting the virial theorem to estimate σ from the same stellar masses by setting $M = c \sigma^2 r_e / G$, we find $\sigma = 300$ – 350 km s $^{-1}$ for galaxies with $M_{\text{stellar}} = 2$ – $4 \times 10^{11} M_\odot$ (see also Nesvadba et al. 2011b). c is a constant that depends on the radial mass profile. We use $c = 5$ (Bender et al. 1992), and adopt $r_e = 2$ – 3 kpc for the effective radius of the galaxy. In Fig. 5 we show where the bolometric AGN luminosity of our galaxies falls relative to the stellar mass and the critical luminosity implied by Murray et al. (2005, red dashed line). All but seven sources have AGN that would not be luminous enough to exceed the critical luminosity to unbind the gas.

7. Radio jets

Centimeter radio jets are powered by the synchrotron emission from relativistic particles, with an emissivity that depends on the

surrounding gas density and magnetic field strength, and is difficult to quantify accurately from first principles. Therefore, a number of empirical measures have been developed to estimate the kinetic power of a radio jet from the observed monochromatic radio power at a given frequency. To estimate the global impact of the radio jets on the gas, we rely upon the relationship of Cavagnolo et al. (2010), who measured the mechanical work done by radio jets in massive galaxy clusters to inflate cavities in the intracluster medium, thereby providing empirical estimates of the work done by the jet against the surrounding gas.

The radio power of the entire SINFONI sample, including the 24 galaxies discussed here, has already been derived by N16 by interpolating the observed multifrequency radio fluxes in the NASA/IPAC Extragalactic Database NED. Results for individual targets are given in their Table 3. We summarize here only the global results for our present targets in the form of the green histograms in Fig. 4. The left panel shows the distribution of the radio power at rest-frame 1.4 GHz. To ease the comparison with the far-infrared luminosity from star formation and AGN, we also measure the radio luminosities in solar units, finding that our sources fall into a range of $\log P_{1.4} = 10^{10} - 10^{12} L_{\odot}$ (which corresponds to $10^{27-29} \text{ W Hz}^{-1}$ at 1.4 GHz in the rest-frame). Our sources thus include some of the most powerful radio sources known at their redshift. Nesvadba et al. (2016) also demonstrate that these estimates are in good agreement with the empirical calibration of, e.g., Willott et al. (2000), and also detailed studies of the X-ray and radio spectral properties in two nearby isolated radio galaxies (Harwood et al. 2016). They are conservative estimates, in that they predict a jet kinetic power that is about 0.2–0.3 dex lower than estimates by Turner & Shabala (2015). The latter take into account that the radio emission should become fainter as the radio lobes expand (see also Kaiser & Alexander 1997).

The right panel of Fig. 4 shows the kinetic power of the radio jets obtained with the Cavagnolo et al. (2010) approach. The values are about an order of magnitude greater than those of the bolometric quasar power, and about three orders of magnitude greater than the kinetic power of the star formation.

In the lower left panel of Fig. 3 we compare the kinetic energy of the gas with that provided by the radio jet. The solid red and blue dashed lines in the right panel show energy conversion efficiencies between 0.1 and 100%. The jets provide sufficient kinetic energy to explain the gas kinematics in all sources with a wide range of efficiencies. Since this energy comes from an extended radio jet, it is being deposited locally into the interstellar gas, which should make the energy transfer particularly efficient (e.g., Mukherjee et al. 2016).

To estimate the momentum carried by the radio jet, we use the calibration of Cavagnolo et al. (2010) already discussed above, and use $p_{\text{jet}} = E_{\text{jet}}/v$, with $v = \beta c$. Observations of compact steep spectrum radio galaxies in the nearby Universe, which are good candidates for being in an early phase of jet activity, where the jet interacts strongly with the dense circumnuclear gas, suggest $\beta = 0.01 - 0.1$ (e.g., Readhead et al. 1996; Owsianik & Conway 1998; Harwood et al. 2013). Most suitable for our purposes is to provide firm lower limits to the momentum injection from the jets, and we therefore adopt $\beta = 0.1$. Keeping in mind that the observed velocities of entrained clouds are lower than the expansion velocity of the hot cocoon gas by at least factors of a few (Cooper et al. 2008; Gaibler et al. 2009), and that we observe velocities of up to $\sim 2500 \text{ km s}^{-1}$ for the warm ionized clouds in our SINFONI sample, jet expansion velocities much smaller than $0.01c$ do not seem very likely, unless the medium is very dense or clumpy, so that the jet finds it difficult to

escape. Hydrodynamic models, however, show, that during such phases jets are particularly efficient in imparting their momentum into the gas. As the jets push through low-density channels in such galaxies, their ram pressure produces mechanical advantages of factors of 10–100, so that the velocities of the entrained gas may reach velocities of a few $10^{2-3} \text{ km s}^{-1}$, as observed (N16), although the jet expansion velocity itself may actually be much lower (Wagner & Bicknell 2011; Wagner et al. 2012). As a result, the galaxies would fall into a very similar range in Fig. 3, although the physics behind the gas acceleration would be more complicated.

The momentum of the gas as a function of the momentum provided by the radio jet in our sources is shown in the bottom right panel of Fig. 3. The black dashed region in Fig. 3 shows the range in the diagram where a jet expansion velocity of less than $0.01c$ would be required, which we consider unlikely. All galaxies fall below this range, i.e., within a region of the diagram where the momentum of the gas can be plausibly powered by the radio source. We note that galaxies with greater jet expansion velocity transfer less of their momentum into the gas, so that it is not unphysical to find galaxies which fall formally in the range where $v_{\text{jet}} > c$. In this case, additional effects, perhaps due to geometry or clumpiness of the gas, could make the momentum transfer less efficient than implied by our simplified assumptions.

8. Ambiguous cases

Our sample includes six targets, for which the simple energy and momentum considerations presented above cannot distinguish between AGN bolometric and radio power as the main driver of the gas kinematics. Both radio source and quasar radiation seem powerful enough to explain their observed gas kinematics (either kinetic energy, momentum, or both), even when adopting the relatively low energy and momentum transfer efficiencies suggested by FeLoBAL observations and the most detailed hydrodynamic models (Moe et al. 2009; Dunn et al. 2010; Bautista et al. 2010; Krumholz & Thompson 2012; Bieri et al. 2017).

These six targets are among those for which we measured the lowest kinetic energy and momentum in our SINFONI data sets. Three of these targets, MRC 2224–273, MRC 0114–211, and TN J2007–1316, are not spatially well resolved, so that beam smearing might lower their velocity measurements and thus their kinetic energy and momentum estimates. This is in particular the case for TN J2007–1316, which has one of the lowest velocity gradients measured in our entire sample, $\Delta v = 100 \text{ km s}^{-1}$ (N16). PKS 0529–549 and 3C 257, in contrast, have well-resolved emission-line regions, and radio jets that are still confined within the extended emission line gas. Both galaxies show kinematic disturbances associated with the extended radio emission, which suggests that most of the energy injection comes from the radio source.

The sixth source, USS 0211–122 has very low line widths in the extended gas, and pronounced blue wings in the [OIII] $\lambda\lambda 4959, 5007$ line profiles (N16). This is unusual compared to the remaining sample of HzRGs with SINFONI data from N16, where line profiles are generally broad, reflecting the large line widths and strong velocity gradients over most of the bright emission-line gas, but do not show pronounced secondary components, as are commonly found in quasars. Collet et al. (2016) propose that this could be an intrinsic difference between radio galaxies and quasars, perhaps indicating different radial sizes of the outflows. If this is correct, then USS 0211–122 could well be a source where the energy injection from the bolometric

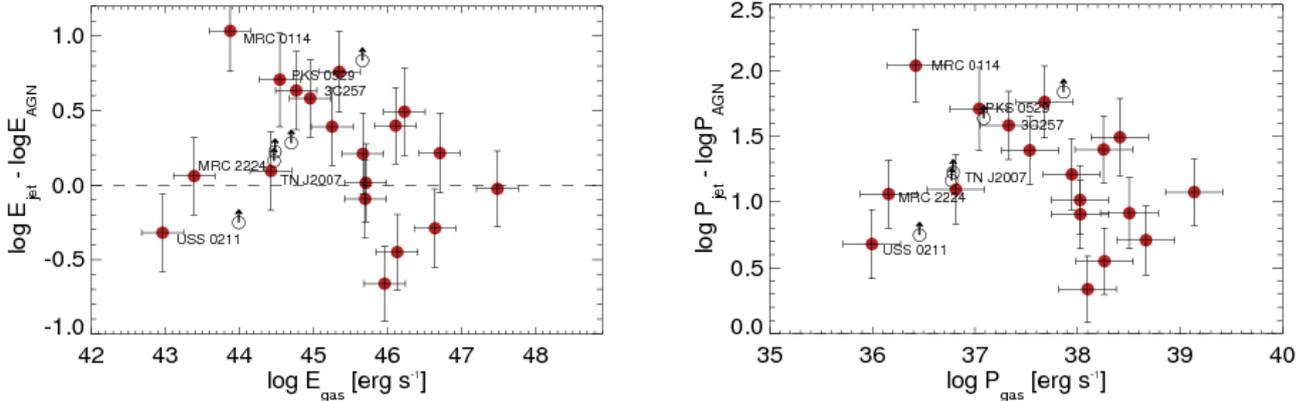


Fig. 6. Ratio of the kinetic energy (*left*) and momentum (*right*) transferred into the gas through AGN radiation and radio jet as a function of those observed in the gas. The dashed line shows where both are equal. We label individual targets, for which a simple comparison of the energy and momentum input rates do not favor either radiation or jets as the main injection mechanism.

quasar radiation exceeds that from the radio jet. This could also explain why surveys of large samples of typically unresolved quasar spectra find that broad, blueshifted wings appear more associated with radio-quiet than radio-loud AGN (Heckman et al. 1984; Greene & Ho 2005; Woo et al. 2016), whereas detailed observations of individual galaxies with radio-loud AGN find perturbed kinematic features preferentially along the radio jet axis (e.g., Husemann et al. 2013).

In Fig. 6 we show the ratio between the energy and momentum injection from radiation and jets as a function of gas kinetic energy. The figure shows that the expected energy and momentum injection related to the jet is typically higher by at least factors of a few than those from the AGN radiation. MRC 0114–211, PKS 0529–549 and 3C 257, are, in fact, among the galaxies, where the power of the radio jet is particularly strong compared to the bolometric power. It would not be obvious why the gas kinematics in these sources should be dominated by the radiation rather than the jet if the opposite were the case in most of the other sources, as suggested by Fig. 3. In USS 0211–122, however, the source with particularly high rest-frame UV polarization (Vernet et al. 2001) already discussed in Sect. 5, the bolometric power is greater than the radio power, and therefore the mechanism powering the gas kinematics remains more ambiguous.

9. Discussion and conclusions

We used observations of the kinematics of warm ionized gas, and the far-infrared and centimeter radio spectral energy distribution of 24 powerful radio galaxies at $z \sim 2$ to compare the rates of kinetic energy released by star formation, quasar radiation, and radio jets with the energy input rates required to power the observed gas kinematics. All galaxies have far-infrared spectral energy distributions measured with *Herschel*/PACS and SPIRE photometry, which Drouart et al. (2014) decomposed into an AGN and star formation component (decompositions for individual targets are shown in their Fig. D.1). They have also been observed with SINFONI to infer the kinematics of the warm ionized gas (Nesvadba et al. 2016). We use standard techniques from the literature to estimate the kinetic energy and momentum released by star formation, radiation from the buried AGN, and radio jets, and to make a comparison with the kinetic energy and momentum found in the gas. We also take into account that the efficiency of the energy and momentum transfer from each source to the gas is unlikely to be unity, and compare this with

several assumptions given in the literature. Our understanding of AGN feedback is still evolving very rapidly, and our goal was to capture the main streams of the current discussion in our analysis while also providing observational constraints that are generic enough to be useful for comparison with future theoretical work.

We find that the kinetic energy produced by starburst-driven winds falls short of what is required to drive the observed gas kinematics in agreement with previous results for individual sources (Nesvadba et al. 2006; Seymour et al. 2012). For radiative AGN feedback (“quasar feedback”) the potential of producing the observed gas kinematics is more ambiguous and depends on the assumed efficiency with which the gas is being accelerated. For the most optimistic assumptions, AGN radiation could be behind the gas kinematics in about two-thirds of our sources. For models including more explicit descriptions of the radiative transfer, lower efficiencies are found. In this case, the quasar emission would generally not be sufficient to power the kinematics of the warm ionized gas. In six galaxies with low gas kinetic energy and momentum, AGN radiation could be enough to power the kinematics even with these relatively low injection rates; however, the specific properties of these sources suggest nonetheless that the radio jet is probably more effective. One notable exception is USS 0211–122, which has unusual [OIII] line profiles more reminiscent of quasars than radio galaxies, and which also has a low expected rate of jet-to-AGN energy and momentum.

Radiation pressure also seems insufficient to unbind the gas from galaxies with stellar masses as high as in our targets. The radio jets appear capable of providing sufficient energy and momentum to produce the highly energized, extended emission-line regions we observe in these galaxies.

The analysis presented here is complementary to that of Nesvadba et al. (2016), who analyzed the detailed properties of 49 HzRGs including the 24 we discuss here, which also have the *Herschel* far-infrared photometry necessary for a detailed discussion of the energy produced by AGN and star formation. They found several signatures, e.g., a general co-alignment of radio jet axis and major axis of the extended emission-line regions, and abrupt changes in velocity or line widths associated with the jet direction or radio features, which also suggest that the radio jet is the main source of energy and momentum injection.

We emphasize that our results do not imply that AGN radiation and star formation have no effect at all on their surrounding gas. With an analysis of the global gas properties like the

one done here, we cannot probe further than to identify the most powerful mechanism. Nevertheless, our results do not require us to assume more complex scenarios, e.g., a conspiracy of gas acceleration through radiation and jets, as would be the case if no single mechanism were able to provide sufficient momentum and energy by itself. Detailed, high-resolution studies of radio-loud quasars at low redshift currently paint a mixed picture; some authors have found that radiation dominates (e.g., ?), and others attribute a larger role to the radio jet (e.g., Husemann et al. 2013).

Our results for this particular moment in the evolution of our sources are also independent from timescale or duty cycle considerations because we observe the impact of each source in situ, and at the same moment when they are acting on the same gas reservoir. With jet lifetimes of a few 10^{6-7} yr (e.g., Blundell & Rawlings 1999), and star formation timescales of at least a few 10^8 and perhaps even up to 10^9 yr (Rocca-Volmerange et al. 2013), it is clear that we are witnessing these processes at a peculiar moment in time, most likely when the active formation phase of the host galaxy is nearly terminated (e.g., Nesvadba et al. 2011a; Drouart et al. 2014). While star formation and AGN radiation have probably affected the host evolution prior to this phase, they have not removed the gas we see in these targets now. Our results therefore suggest that the radio jet dominates the final “sweeping clean” of gas from these very massive galaxies. This is also consistent with the result presented in Fig. 5 that AGN with radiative power of 10^{46-47} erg s $^{-1}$, in the range of the most luminous AGN at these redshifts, do not seem able to overcome the gravitational potential of galaxies with $M \geq 10^{11} M_{\odot}$ through the momentum carried by their bolometric radiation alone. In earlier stages of assembly of our galaxies, and lower mass galaxies generally, the relative balance of star formation, AGN radiation, and radio jets, might therefore be very different.

Making this comparison for the present set of galaxies was relatively simple because the sources are well extended and the radio sources are bright enough that their kinetic power dominates globally over other mechanisms. However, as one moves down the radio luminosity function, this becomes more difficult, as can be seen from a very similar analysis performed by Nesvadba et al. (2011b) and Polletta et al. (2011) for two of the brightest obscured quasars from the *Spitzer* SWIRE survey (Polletta 2008). The obscured quasars in these sources are about as powerful as the ones we discuss here, but their radio sources are about four orders of magnitude fainter than the brightest sources we discuss here (Sajina et al. 2007). Nonetheless, their gas kinematics seem to be dominated by the radio jets, contrary to AGN selected, for example, in the optical or X-ray. Much deeper and more detailed analyses with high-resolution radio data will be required to study galaxies where the energy output from star formation, quasar, and jets are more balanced. However, this will be a long-term effort, as currently, no general, reliable, uncontested methods exist to distinguish between star formation and AGN in faint high-redshift radio sources. ALMA, the next generations of ELTs and JWST in the optical and near-infrared, and radio facilities like LOFAR and SKA will undoubtedly play a large role in disentangling the relative contribution of each process for the regulation of the gas kinematics in the general population of high-redshift AGN host galaxies.

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References

- Archibald, E. N., Dunlop, J. S., Hughes, D. H., et al. 2001, *MNRAS*, 323, 417
- Bautista, M. A., Dunn, J. P., Arav, N., et al. 2010, *ApJ*, 713, 25
- Bender, R., Burstein, D., & Faber, S. M. 1992, *ApJ*, 399, 462
- Bieri, R., Dubois, Y., Rosdahl, J., et al. 2017, *MNRAS*, 464, 1854
- Blundell, K. M., & Rawlings, S. 1999, *Nature*, 399, 330
- Buitrago, F., Conselice, C. J., Epinat, B., et al. 2014, *MNRAS*, 439, 1494
- Carilli, C. L., Röttgering, H. J. A., van Ojik, R., Miley, G. K., & van Breugel, W. J. M. 1997, *ApJS*, 109, 1
- Carilli, C. L., Harris, D. E., Pentericci, L., et al. 2002, *ApJ*, 567, 781
- Cavagnolo, K. W., McNamara, B. R., Nulsen, P. E. J., et al. 2010, *ApJ*, 720, 1066
- Chabrier, G. 2003, *PASP*, 115, 763
- Cicone, C., Maiolino, R., Sturm, E., et al. 2014, *A&A*, 562, A21
- Collet, C., Nesvadba, N. P. H., De Breuck, C., et al. 2015, *A&A*, 579, A89
- Collet, C., Nesvadba, N. P. H., De Breuck, C., et al. 2016, *A&A*, 586, A152
- Cooper, J. L., Bicknell, G. V., Sutherland, R. S., & Bland-Hawthorn, J. 2008, *ApJ*, 674, 157
- Dalla Vecchia, C., & Schaye, J. 2008, *MNRAS*, 387, 1431
- De Breuck, C., van Breugel, W., Röttgering, H., et al. 2001, *AJ*, 121, 1241
- De Breuck, C., Seymour, N., Stern, D., et al. 2010, *ApJ*, 725, 36
- Di Matteo, T., Springel, V., & Hernquist, L. 2005, *Nature*, 433, 604
- Drouart, G., De Breuck, C., Vernet, J., et al. 2014, *A&A*, 566, A53
- Drouart, G., Rocca-Volmerange, B., De Breuck, C., et al. 2016, *A&A*, 593, A109
- Dunn, J. P., Bautista, M., Arav, N., et al. 2010, *ApJ*, 709, 611
- Elvis, M., Wilkes, B. J., McDowell, J. C., et al. 1994, *ApJS*, 95, 1
- Emonts, B. H. C., Norris, R. P., Feain, I., et al. 2014, *MNRAS*, 438, 2898
- Faucher-Giguère, C.-A., & Quataert, E. 2012, *MNRAS*, 425, 605
- Gaibler, V., Krause, M., & Camenzind, M. 2009, *MNRAS*, 400, 1785
- Greene, J. E., & Ho, L. C. 2005, *ApJ*, 630, 122
- Gullberg, B., De Breuck, C., Lehnert, M. D., et al. 2016, *A&A*, 586, A124
- Harwood, J. J., Hardcastle, M. J., Croston, J. H., & Goodger, J. L. 2013, *MNRAS*, 435, 3353
- Harwood, J. J., Croston, J. H., Intema, H. T., et al. 2016, *MNRAS*, 458, 4443
- Hatch, N. A., Overzier, R. A., Röttgering, H. J. A., Kurk, J. D., & Miley, G. K. 2008, *MNRAS*, 383, 931
- Hatch, N. A., Overzier, R. A., Kurk, J. D., et al. 2009, *MNRAS*, 395, 114
- Hatch, N. A., Röttgering, H. J. A., Miley, G. K., et al. 2013, *MNRAS*, 436, 2244
- Heckman, T. M., Miley, G. K., & Green, R. F. 1984, *ApJ*, 281, 525
- Heckman, T. M., Armus, L., & Miley, G. K. 1990, *ApJS*, 74, 833
- Heckman, T. M., Alexandroff, R. M., Borthakur, S., Overzier, R., & Leitherer, C. 2015, *ApJ*, 809, 147
- Humphrey, A., Iwamuro, F., Villar-Martín, M., Binette, L., & Sung, E. C. 2009, *MNRAS*, 399, L34
- Husemann, B., Wisotzki, L., Sánchez, S. F., & Jahnke, K. 2013, *A&A*, 549, A43
- Ivison, R. J., Morrison, G. E., Biggs, A. D., et al. 2008, *MNRAS*, 390, 1117
- Kaiser, C. R., & Alexander, P. 1997, *MNRAS*, 286, 215
- Kennicutt, Jr., R. C. 1998, *ARA&A*, 36, 189
- Krumholz, M. R., & Thompson, T. A. 2012, *ApJ*, 760, 155
- Krumholz, M. R., & Thompson, T. A. 2013, *MNRAS*, 434, 2329
- Lehnert, M. D., & Heckman, T. M. 1996a, *ApJ*, 462, 651
- Lehnert, M. D., & Heckman, T. M. 1996b, *ApJ*, 472, 546
- Liu, G., Zakamska, N. L., Greene, J. E., Nesvadba, N. P. H., & Liu, X. 2013a, *MNRAS*, 430, 2327
- Liu, G., Zakamska, N. L., Greene, J. E., Nesvadba, N. P. H., & Liu, X. 2013b, *MNRAS*, 436, 2576
- McCarthy, P. J. 1999, in *The Most Distant Radio Galaxies*, eds. H. J. A. Röttgering, P. N. Best, & M. D. Lehnert, 5
- Moe, M., Arav, N., Bautista, M. A., & Korista, K. T. 2009, *ApJ*, 706, 525
- Mukherjee, D., Bicknell, G. V., Sutherland, R., & Wagner, A. 2016, *MNRAS*, 461, 967
- Mullaney, J. R., Alexander, D. M., Goulding, A. D., & Hickox, R. C. 2011, *MNRAS*, 414, 1082
- Murray, N., Quataert, E., & Thompson, T. A. 2005, *ApJ*, 618, 569

- Nesvadba, N., De Breuck, C., et al. 2016, *A&A*, accepted
- Nesvadba, N. P. H., Lehnert, M. D., Eisenhauer, F., et al. 2006, *ApJ*, 650, 693
- Nesvadba, N. P. H., Lehnert, M. D., De Breuck, C., Gilbert, A., & van Breugel, W. 2007, *A&A*, 475, 145
- Nesvadba, N. P. H., Lehnert, M. D., Davies, R. I., Verma, A., & Eisenhauer, F. 2008a, *A&A*, 479, 67
- Nesvadba, N. P. H., Lehnert, M. D., De Breuck, C., Gilbert, A. M., & van Breugel, W. 2008b, *A&A*, 491, 407
- Nesvadba, N. P. H., Neri, R., De Breuck, C., et al. 2009, *MNRAS*, 395, L16
- Nesvadba, N. P. H., De Breuck, C., Lehnert, M. D., et al. 2011a, *A&A*, 525, A43
- Nesvadba, N. P. H., Polletta, M., Lehnert, M. D., et al. 2011b, *MNRAS*, 415, 2359
- Novak, G. S., Ostriker, J. P., & Ciotti, L. 2012, *MNRAS*, 427, 2734
- Overzier, R. A., Harris, D. E., Carilli, C. L., et al. 2005, *A&A*, 433, 87
- Owsianik, I., & Conway, J. E. 1998, *A&A*, 337, 69
- Pentericci, L., Van Reeve, W., Carilli, C. L., Röttgering, H. J. A., & Miley, G. K. 2000, *A&AS*, 145, 121
- Polletta, M. 2008, *A&A*, 480, L41
- Polletta, M., Nesvadba, N. P. H., Neri, R., et al. 2011, *A&A*, 533, A20
- Readhead, A. C. S., Taylor, G. B., Pearson, T. J., & Wilkinson, P. N. 1996, *ApJ*, 460, 634
- Reuland, M., Röttgering, H., van Breugel, W., & De Breuck, C. 2004, *MNRAS*, 353, 377
- Richards, G. T., Lacy, M., Storrie-Lombardi, L. J., et al. 2006, *ApJS*, 166, 470
- Rocca-Volmerange, B., Drouart, G., De Breuck, C., et al. 2013, *MNRAS*, 429, 2780
- Sajina, A., Yan, L., Lacy, M., & Huynh, M. 2007, *ApJ*, 667, L17
- Seymour, N., Stern, D., De Breuck, C., et al. 2007, *ApJS*, 171, 353
- Seymour, N., Altieri, B., De Breuck, C., et al. 2012, *ApJ*, 755, 146
- Sharma, P., Roy, A., Nath, B. B., & Shchekinov, Y. 2014, *MNRAS*, 443, 3463
- Silk, J., & Rees, M. J. 1998, *A&A*, 331, L1
- Strickland, D. K., & Heckman, T. M. 2009, *ApJ*, 697, 2030
- Sun, A.-L., Greene, J. E., Zakamska, N. L., & Nesvadba, N. P. H. 2014, *ApJ*, 790, 160
- Tadhunter, C., Dicken, D., Morganti, R., et al. 2014, *MNRAS*, 445, L51
- Taniguchi, Y., Ohyama, Y., Murayama, T., et al. 2001, *ApJ*, 559, L9
- Turner, R. J., & Shabala, S. S. 2015, *ApJ*, 806, 59
- Veilleux, S., Cecil, G., & Bland-Hawthorn, J. 2005, *ARA&A*, 43, 769
- Veilleux, S., Meléndez, M., Sturm, E., et al. 2013, *ApJ*, 776, 27
- Vernet, J., Fosbury, R. A. E., Villar-Martín, M., et al. 2001, *A&A*, 366, 7
- Villar-Martín, M., Vernet, J., di Serego Alighieri, S., et al. 2003, *MNRAS*, 346, 273
- Wagner, A. Y., & Bicknell, G. V. 2011, *ApJ*, 738, 117
- Wagner, A. Y., Bicknell, G. V., & Umemura, M. 2012, *ApJ*, 757, 136
- Willott, C. J., Rawlings, S., & Jarvis, M. J. 2000, *MNRAS*, 313, 237
- Woo, J.-H., Bae, H.-J., Son, D., & Karouzos, M. 2016, *ApJ*, 817, 108
- Zirm, A. W., Overzier, R. A., Miley, G. K., et al. 2005, *ApJ*, 630, 68
- Zubovas, K., & King, A. 2012, *ApJ*, 745, L34